



Nuclear-recoil energy scale in CDMS II silicon dark-matter detectors

R. Agnese¹, A.J. Anderson², T. Aramaki³, W. Baker⁴, D. Balakishiyeva⁵, S. Banik⁶, D. Barker⁷, R. Basu Thakur^{8,9}, D.A. Bauer⁸, T. Binder¹⁰, A. Borgland³, M.A. Bowles^{11,*}, P.L. Brink³, R. Bunker¹², B. Cabrera¹³, D.O. Caldwell^{14,a}, R. Calkins⁵, C. Cartaro³, D.G. Cerdeño^{15,16}, Y.-Y. Chang¹⁷, H. Chagani⁷, Y. Chen^{18,19}, J. Cooley⁵, B. Cornell¹⁷, P. Cushman⁷, M. Daal²⁰, T. Doughty²⁰, E.M. Dragowsky^{21,22}, L. Esteban¹⁵, S. Fallows^{7,19}, E. Fascione²³, E. Figueroa-Feliciano²⁴, M. Fritts⁷, G. Gerbier²³, R. Germond²³, M. Ghaith²³, G.L. Godfrey³, S.R. Golwala¹⁷, J. Hall²⁵, H.R. Harris⁴, D. Holmgren⁸, Z. Hong²⁴, L. Hsu⁸, M.E. Huber^{26,27}, V. Iyer⁶, D. Jardin⁵, A. Jastram⁴, C. Jena⁶, M.H. Kelsey³, A. Kennedy⁷, A. Kubik⁴, N.A. Kurinsky³, A. Leder², E. Lopez Asamar¹⁶, P. Lukens⁸, D. MacDonell^{28,29}, R. Mahapatra⁴, V. Mandic⁷, N. Mast⁷, K.A. McCarthy³⁰, E.H. Miller¹¹, N. Mirabolfathi⁴, R.A. Moffatt¹³, B. Mohanty⁶, D. Moore^{17,31}, J.D. Morales Mendoza⁴, J. Nelson⁷, S.M. Oser^{28,29}, K. Page²³, W.A. Page^{28,29}, R. Partridge³, M. Penalver Martinez¹⁶, M. Pepin⁷, A. Phipps²⁰, S. Poudel¹⁰, M. Pyle²⁰, H. Qiu⁵, W. Rau²³, P. Redl¹³, A. Reissetter³², A. Roberts²⁶, H.E. Rogers⁷, A.E. Robinson³³, T. Saab¹, B. Sadoulet^{20,34}, J. Sander¹⁰, K. Schneck³, R.W. Schnee¹¹, S. Scorza²⁵, K. Senapati⁶, B. Serfass²⁰, D. Speller²⁰, P.C.F. Di Stefano²³, M. Stein⁵, J. Street¹¹, H.A. Tanaka³⁵, D. Toback⁴, R. Underwood²³, A.N. Villano^{7,26}, B. von Krosigk^{28,29}, B. Welliver¹, J.S. Wilson⁴, M.J. Wilson³⁵, D.H. Wright³, S. Yellin¹³, J.J. Yen¹³, B.A. Young³⁶, X. Zhang²³, X. Zhao⁴

¹ Department of Physics, University of Florida, Gainesville, FL 32611, USA

² Department of Physics, Massachusetts Institute of Technology, Cambridge, MA 02139, USA

³ SLAC National Accelerator Laboratory/Kavli Institute for Particle Astrophysics and Cosmology, Menlo Park, CA 94025, USA

⁴ Department of Physics and Astronomy, and the Mitchell Institute for Fundamental Physics and Astronomy, Texas A&M University, College Station, TX 77843, USA

⁵ Department of Physics, Southern Methodist University, Dallas, TX 75275, USA

⁶ School of Physical Sciences, National Institute of Science Education and Research, HBNI, Jatni 752050, India

⁷ School of Physics & Astronomy, University of Minnesota, Minneapolis, MN 55455, USA

⁸ Fermi National Accelerator Laboratory, Batavia, IL 60510, USA

⁹ Department of Physics, University of Illinois at Urbana-Champaign, Urbana, IL 61801, USA

¹⁰ Department of Physics, University of South Dakota, Vermillion, SD 57069, USA

¹¹ Department of Physics, South Dakota School of Mines & Technology, Rapid City, SD 57701, USA

¹² Pacific Northwest National Laboratory, Richland, WA 99352, USA

¹³ Department of Physics, Stanford University, Stanford, CA 94305, USA

¹⁴ Department of Physics, University of California, Santa Barbara, CA 93106, USA

¹⁵ Instituto de Física Teórica UAM/CSIC, Universidad Autónoma de Madrid, 28049 Madrid, Spain

¹⁶ Department of Physics, Durham University, Durham DH1 3LE, UK

¹⁷ Division of Physics, Mathematics, & Astronomy, California Institute of Technology, Pasadena, CA 91125, USA

¹⁸ Department of Physics, Syracuse University, Syracuse, NY 13244, USA

¹⁹ Department of Physics, University of Alberta, Edmonton, T6G 2E1, Canada

²⁰ Department of Physics, University of California, Berkeley, CA 94720, USA

²¹ Department of Physics, Case Western Reserve University, Cleveland, OH 44106, USA

²² Research Computing, University Technologies, Case Western Reserve University, Cleveland, OH 44106, USA

²³ Department of Physics, Queen's University, Kingston, ON K7L 3N6, Canada

²⁴ Department of Physics & Astronomy, Northwestern University, Evanston, IL 60208-3112, USA

²⁵ SNOLAB, Creighton Mine #9, 1039 Regional Road 24, Sudbury, ON P3Y 1N2, Canada

²⁶ Department of Physics, University of Colorado Denver, Denver, CO 80217, USA

* Corresponding author.

E-mail address: Michael.Bowles@Mines.sdsmt.edu (M.A. Bowles).

^a Deceased.

²⁷ Departments of Physics and Electrical Engineering, University of Colorado Denver, Denver, CO 80217, USA

²⁸ Department of Physics & Astronomy, University of British Columbia, Vancouver, BC V6T 1Z1, Canada

²⁹ TRIUMF, Vancouver, BC V6T 2A3, Canada

³⁰ Institute for Disease Modeling, 3150 139th Ave SE, Bellevue, WA 98005, USA

³¹ Department of Physics, Yale University, New Haven, CT 06520, USA

³² Department of Physics, University of Evansville, Evansville, IN 47722, USA

³³ Département de Physique, Université de Montréal, Montréal, Québec H3T 1J4, Canada

³⁴ Lawrence Berkeley National Laboratory, Berkeley, CA 94720, USA

³⁵ Department of Physics, University of Toronto, Toronto, ON M5S 1A7, Canada

³⁶ Department of Physics, Santa Clara University, Santa Clara, CA 95053, USA

ARTICLE INFO

Keywords:

Dark matter

Detector calibration

Nuclear-recoil energy scale

Ionization yield

ABSTRACT

The Cryogenic Dark Matter Search (CDMS II) experiment aims to detect dark matter particles that elastically scatter from nuclei in semiconductor detectors. The resulting nuclear-recoil energy depositions are detected by ionization and phonon sensors. Neutrons produce a similar spectrum of low-energy nuclear recoils in such detectors, while most other backgrounds produce electron recoils. The absolute energy scale for nuclear recoils is necessary to interpret results correctly. The energy scale can be determined in CDMS II silicon detectors using neutrons incident from a broad-spectrum ^{252}Cf source, taking advantage of a prominent resonance in the neutron elastic scattering cross section of silicon at a recoil (neutron) energy near 20 (182) keV. Results indicate that the phonon collection efficiency for nuclear recoils is $4.8^{+0.7}_{-0.9}\%$ lower than for electron recoils of the same energy. Comparisons of the ionization signals for nuclear recoils to those measured previously by other groups at higher electric fields indicate that the ionization collection efficiency for CDMS II silicon detectors operated at ~ 4 V/cm is consistent with 100% for nuclear recoils below 20 keV and gradually decreases for larger energies to $\sim 75\%$ at 100 keV. The impact of these measurements on previously published CDMS II silicon results is small.

1. Introduction

Strong evidence indicates that $\geq 80\%$ of the matter in the Universe is non-luminous and non-baryonic [1]. Weakly Interacting Massive Particles (WIMPs) are a leading class of theoretically motivated candidates for this dark matter [2]. These particles are expected to interact with normal matter through the weak nuclear force and to cluster gravitationally. If WIMPs do constitute our galaxy's dark matter, they may be detectable through their elastic scattering off atomic nuclei in terrestrial detectors [3]. Under standard galactic halo assumptions [4] for a WIMP mass of $\sim 100 \text{ GeV}/c^2$, the recoiling nuclei have energies of tens of keV and ranges of 10–100 nm in solid matter.

The Cryogenic Dark Matter Search (CDMS II) experiment measured nuclear recoils using a target mass composed of high-purity silicon and germanium semiconductor crystals operated at ~ 50 mK. Each crystal was instrumented to simultaneously measure the electron–hole pairs (ionization) and athermal phonons created by particle interactions within the crystal [5].

A WIMP, or a neutron, may scatter off a nucleus producing a nuclear recoil (NR), while most other interactions produce an electron recoil (ER). Accurate determination of an event's energy requires a systematic calibration of the recoil energy scale. This energy calibration is generally straightforward for electron recoils due to the availability of a variety of spectral lines from radioactive sources over a wide range of energies.

The calibration for nuclear recoils is more difficult. CDMS II used a ^{252}Cf neutron source to perform nuclear-recoil calibrations, and the spectrum of recoil energies in CDMS II detectors resulting from exposure to this source decreases quasi-exponentially with increasing energy and is nearly featureless. For CDMS II detectors, knowledge of the nuclear-recoil energy scale to within $\sim 10\%$ is sufficient to accurately interpret WIMP-search results for WIMP masses greater than a few tens of GeV/c^2 . For lower masses, however, a more accurate determination of the energy scale is important for a robust comparison of results from different experiments, particularly in light of interpretations of data from several experiments as possible evidence for a low-mass ($< 10 \text{ GeV}/c^2$) WIMP [6–9].

This paper describes the procedure used to calibrate the nuclear recoil response of CDMS II silicon detectors. Experimental data for this study are drawn from the final runs of these detectors at the

Soudan Underground Laboratory, from July 2007 to September 2008, as described in Ref. [9]. *In situ* measurements of elastic neutron scatters in these detectors from a ^{252}Cf source are compared to Monte Carlo simulations of recoiling nuclei in the detectors. A re-calibrated energy scale is derived, optimizing agreement between measured and simulated recoil spectra. This is used to adjust the published upper limits on the WIMP-nucleon spin-independent cross section, as well as the 90% C.L. acceptance region from the analysis of the final exposure of the silicon detectors [9].

2. CDMS II detectors

The final configuration of CDMS II contained 11 silicon and 19 germanium Z-sensitive Ionization- and Phonon-mediated (ZIP) detectors. These were arrayed into five “towers”, each containing six detectors following the designation $\text{T}x\text{Z}y$ where x (1–5) is the tower number and y (1–6) indicates the position within the stack (from top to bottom). We focus here on the silicon detectors used in Ref. [9], which were ~ 10 mm thick, 76 mm in diameter, with a mass of ~ 106 g each. Of the eleven silicon detectors, two were excluded due to wiring failures leading to incomplete ionization collection, and a third was excluded due to unstable phonon channel response.

Each detector was photolithographically patterned with sensors on both flat faces: two concentric ionization electrodes on one face and four independent phonon sensors on the opposite face. The ionization electrodes were biased to 4 V with respect to the phonon electrodes, creating an electric field of 4 V/cm in the bulk of the detector along its z axis [10]. The electrons and holes generated by a particle interaction were separated and drifted across the crystal by the electric field, generating image currents in the electrodes detected by a JFET-based charge amplifier [11]. By careful neutralization of ionized trapping sites within the crystal with regular exposure to infrared LEDs (“flashing”), the detectors were operated in a metastable state in which trapping of charge carriers in the crystal bulk was low. The ionization collection efficiency for electron recoils was therefore high, despite the relatively modest applied electric field.

In semiconductor devices such as the ZIPs, phonon (φ) energy is generated by three interactions: the initial recoil generates primary phonons, the work done on the charge carriers by the electric field generates Neganov–Trofimov–Luke (or NTL) phonons [12–14], and charge carrier relaxation to the Fermi level at the electrodes generates

Download English Version:

<https://daneshyari.com/en/article/8165897>

Download Persian Version:

<https://daneshyari.com/article/8165897>

[Daneshyari.com](https://daneshyari.com)